

Inflation with light Weyl ghost

A. Tokareva (INR, Moscow) in collaboration with M. Ivanov (EPFL, Lausanne), D. Gorbunov (INR, Moscow)

May, 10

- 1 Motivation
- 2 How to hide a ghost
- 3 Inflationary perturbations
- 4 Results for the ghost perturbations
- 5 Conclusions

The renormalizable quadratic gravity

Action:

$$S = \frac{M_P^2}{2} \int d^4x \sqrt{-g} \left(R - \frac{\gamma}{2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} + \frac{\beta}{6} R^2 \right)$$

The renormalizable quadratic gravity

Action:

$$S = \frac{M_P^2}{2} \int d^4x \sqrt{-g} \left(R - \frac{\gamma}{2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} + \frac{\beta}{6} R^2 \right)$$



renormalizable but non-unitary gravity (K.S.Stelle, 1977)

The renormalizable quadratic gravity

Action:

$$S = \frac{M_p^2}{2} \int d^4x \sqrt{-g} \left(R - \frac{\gamma}{2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} + \frac{\beta}{6} R^2 \right)$$

⇓

renormalizable but non-unitary gravity (K.S.Stelle, 1977)

Recently proposed adimensional gravity (agravity): (A.Salvio and A.Strumia, 1403.4226)

$$S = \int d^4x \sqrt{-g} \left(-\frac{1}{f_2^2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} + \frac{1}{6f_0^2} R^2 - \right. \\ \left. - |\partial_\mu S|^2 - \lambda_S |S|^4 + \lambda_{HS} |S|^2 |H|^2 + \xi_S |S|^2 R \right)$$

The renormalizable quadratic gravity

Action:

$$S = \frac{M_p^2}{2} \int d^4x \sqrt{-g} \left(R - \frac{\gamma}{2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} + \frac{\beta}{6} R^2 \right)$$

⇓

renormalizable but non-unitary gravity (K.S.Stelle, 1977)

Recently proposed adimensional gravity (agravity): (A.Salvio and A.Strumia, 1403.4226)

$$S = \int d^4x \sqrt{-g} \left(-\frac{1}{f_2^2} C_{\mu\nu\rho\sigma} C^{\mu\nu\rho\sigma} + \frac{1}{6f_0^2} R^2 - \right. \\ \left. - |\partial_\mu S|^2 - \lambda_S |S|^4 + \lambda_{HS} |S|^2 |H|^2 + \xi_S |S|^2 R \right)$$

In both frameworks there is massive spin-2 ghost.

Effective non-unitary theory?

- Usually if you deal with the non-renormalizable theory you use the effective field theory to control quantum corrections

Effective non-unitary theory?

- Usually if you deal with the non-renormalizable theory you use the effective field theory to control quantum corrections
- **Example:** For Higgs inflation it's impossible to connect inflation with low energy physics without UV completion. (Sibiryakov et. al., arXiv:1008.5157)

Effective non-unitary theory?

- Usually if you deal with the non-renormalizable theory you use the effective field theory to control quantum corrections
- **Example:** For Higgs inflation it's impossible to connect inflation with low energy physics without UV completion. (Sibiryakov et. al., arXiv:1008.5157)
- **The main proposal:** use the renormalizable gravity as UV completion and try to hide the non-unitarity (if possible).

Effective non-unitary theory?

- Usually if you deal with the non-renormalizable theory you use the effective field theory to control quantum corrections
- **Example:** For Higgs inflation it's impossible to connect inflation with low energy physics without UV completion. (Sibiryakov et. al., arXiv:1008.5157)
- **The main proposal:** use the renormalizable gravity as UV completion and try to hide the non-unitarity (if possible).
- Lets quantize the ghost with negative norm \Rightarrow **positive energy, no infinite integrals**, but **negative probabilities**.

Effective non-unitary theory?

- Usually if you deal with the non-renormalizable theory you use the effective field theory to control quantum corrections
- **Example:** For Higgs inflation it's impossible to connect inflation with low energy physics without UV completion. (Sibiryakov et. al., arXiv:1008.5157)
- **The main proposal:** use the renormalizable gravity as UV completion and try to hide the non-unitarity (if possible).
- Lets quantize the ghost with negative norm \Rightarrow **positive energy, no infinite integrals**, but **negative probabilities**.
- **What is the ghost impact on inflation?**

Effective non-unitary theory?

- Usually if you deal with the non-renormalizable theory you use the effective field theory to control quantum corrections
- **Example:** For Higgs inflation it's impossible to connect inflation with low energy physics without UV completion. (Sibiryakov et. al., arXiv:1008.5157)
- **The main proposal:** use the renormalizable gravity as UV completion and try to hide the non-unitarity (if possible).
- Lets quantize the ghost with negative norm \Rightarrow **positive energy, no infinite integrals**, but **negative probabilities**.
- **What is the ghost impact on inflation?**
 - Heavy ghost, $m > H_i$: studied in Sasaki et al. arXiv:0907.3868, arXiv:1012.5202.
Small corrections to the amplitude of tensor perturbations and no corrections for the scalar perturbations. No ghost production.

Effective non-unitary theory?

- Usually if you deal with the non-renormalizable theory you use the effective field theory to control quantum corrections
- **Example:** For Higgs inflation it's impossible to connect inflation with low energy physics without UV completion. (Sibiryakov et. al., arXiv:1008.5157)
- **The main proposal:** use the renormalizable gravity as UV completion and try to hide the non-unitarity (if possible).
- Lets quantize the ghost with negative norm \Rightarrow **positive energy, no infinite integrals**, but **negative probabilities**.
- **What is the ghost impact on inflation?**
 - Heavy ghost, $m > H_i$: studied in Sasaki et al. arXiv:0907.3868, arXiv:1012.5202. Small corrections to the amplitude of tensor perturbations and no corrections for the scalar perturbations. No ghost production.
 - Light ghost, $m < H_i$: **amplitude of tensor perturbations damped by factor $m/\sqrt{2}H_i$** . We studied the scalar perturbations for that case.

May the light ghost be safe?

- For $m \gtrsim 10^{10}$ GeV – the **naturalness problem** for the Higgs mass. Ghost impact on the ratio m_H/M_P becomes large (A.Salvio and A.Strumia, 1403.4226).

May the light ghost be safe?

- For $m \gtrsim 10^{10}$ GeV – the **naturalness problem** for the Higgs mass. Ghost impact on the ratio m_H/M_P becomes large (A.Salvio and A.Strumia, 1403.4226).
- The great problem of interpretation exists only for the **interacting** ghost. The coupling with matter is of order p/M_P , the decay rate is $\Gamma \sim m^3/M_P^2$.

May the light ghost be safe?

- For $m \gtrsim 10^{10}$ GeV – the **naturalness problem** for the Higgs mass. Ghost impact on the ratio m_H/M_P becomes large (A.Salvio and A.Strumia, 1403.4226).
- The great problem of interpretation exists only for the **interacting** ghost. The coupling with matter is of order p/M_P , the decay rate is $\Gamma \sim m^3/M_P^2$. For $m \lesssim 10\text{MeV}$ ghost **lives longer than Universe**.

May the light ghost be safe?

- For $m \gtrsim 10^{10}$ GeV – the **naturalness problem** for the Higgs mass. Ghost impact on the ratio m_H/M_P becomes large (A.Salvio and A.Strumia, 1403.4226).
- The great problem of interpretation exists only for the **interacting** ghost. The coupling with matter is of order p/M_P , the decay rate is $\Gamma \sim m^3/M_P^2$. For $m \lesssim 10\text{MeV}$ ghost **lives longer than Universe**.
- Ghost production in the hot plasma:

$$\sigma \sim e^2/M_P^2,$$

May the light ghost be safe?

- For $m \gtrsim 10^{10}$ GeV – the **naturalness problem** for the Higgs mass. Ghost impact on the ratio m_H/M_P becomes large (A.Salvio and A.Strumia, 1403.4226).
- The great problem of interpretation exists only for the **interacting** ghost. The coupling with matter is of order p/M_P , the decay rate is $\Gamma \sim m^3/M_P^2$. For $m \lesssim 10 \text{ MeV}$ ghost **lives longer than Universe**.
- Ghost production in the hot plasma:

$$\sigma \sim e^2/M_P^2, \quad \Omega_{gh} = 0.02 \frac{m}{10 \text{ MeV}} \frac{T_{max}}{10^{13} \text{ GeV}}$$

May the light ghost be safe?

- For $m \gtrsim 10^{10}$ GeV – the **naturalness problem** for the Higgs mass. Ghost impact on the ratio m_H/M_P becomes large (A.Salvio and A.Strumia, 1403.4226).
- The great problem of interpretation exists only for the **interacting** ghost. The coupling with matter is of order p/M_P , the decay rate is $\Gamma \sim m^3/M_P^2$. For $m \lesssim 10 \text{ MeV}$ ghost **lives longer than Universe**.
- Ghost production in the hot plasma:

$$\sigma \sim e^2/M_P^2, \quad \Omega_{gh} = 0.02 \frac{m}{10 \text{ MeV}} \frac{T_{max}}{10^{13} \text{ GeV}}$$

- From short distance constraints: $m > 4 \cdot 10^{-12} \text{ GeV}$ (S. J. Smullin, A. A. Geraci, D. M. Weld, A. Kapitulnik and J. Chiaverini, eConf C 040802, MOT004 (2004).)

Linearized action in the newtonian gauge at inflation

$$ds^2 = a(\eta)^2 \{ -(1+2\Psi) d\eta^2 + 2\bar{\Psi}_i d\eta dx^i + [(1+2\Phi) \delta_{ij} + \bar{h}_{ij}] dx^i dx^j \}, \quad \delta\phi = \chi.$$

Linearized action in the newtonian gauge at inflation

$$ds^2 = a(\eta)^2 \{ -(1+2\Psi) d\eta^2 + 2\bar{\Psi}_i d\eta dx^i + [(1+2\Phi) \delta_{ij} + \bar{h}_{ij}] dx^i dx^j \}, \quad \delta\phi = \chi.$$

$$S^{(S)} = S_E^{(S)} - \frac{\gamma M_P^2}{3} \int d^4x [\Delta(\Psi - \Phi)]^2$$

$$\begin{aligned} \frac{S_E^{(S)}}{M_P^2} = & \frac{1}{2} \int d^4x a^2 \left[-6\Phi'^2 + 12\mathcal{H}\Psi\Phi' + 2\partial_i\Phi(2\partial^i\Psi + \partial^i\Phi) - 2(\mathcal{H}' + 2\mathcal{H}^2)\Psi^2 \right. \\ & \left. + \frac{1}{M_P^2} (\chi'^2 - \partial_i\chi\partial^i\chi - a^2 V_{,\phi\phi}\chi^2 - 6\phi'\Phi'\chi - 2\phi'\chi'\Psi - 2a^2 V_{,\phi}\Psi\chi) \right] \end{aligned}$$

Linearized action in the newtonian gauge at inflation

$$ds^2 = a(\eta)^2 \{ -(1+2\Psi) d\eta^2 + 2\bar{\Psi}_i d\eta dx^i + [(1+2\Phi) \delta_{ij} + \bar{h}_{ij}] dx^i dx^j \}, \quad \delta\phi = \chi.$$

$$S^{(S)} = S_E^{(S)} - \frac{\gamma M_P^2}{3} \int d^4x [\Delta(\Psi - \Phi)]^2$$

$$\begin{aligned} \frac{S_E^{(S)}}{M_P^2} = & \frac{1}{2} \int d^4x a^2 \left[-6\Phi'^2 + 12\mathcal{H}\Psi\Phi' + 2\partial_i\Phi(2\partial^i\Psi + \partial^i\Phi) - 2(\mathcal{H}' + 2\mathcal{H}^2)\Psi^2 \right. \\ & \left. + \frac{1}{M_P^2} (\chi'^2 - \partial_i\chi\partial^i\chi - a^2 V_{,\phi\phi}\chi^2 - 6\phi'\Phi'\chi - 2\phi'\chi'\Psi - 2a^2 V_{,\phi}\Psi\chi) \right] \end{aligned}$$

For de Sitter background – [exact diagonalization](#).

Linearized action in the newtonian gauge at inflation

$$ds^2 = a(\eta)^2 \{ -(1+2\Psi) d\eta^2 + 2\bar{\Psi}_i d\eta dx^i + [(1+2\Phi) \delta_{ij} + \bar{h}_{ij}] dx^i dx^j \}, \quad \delta\phi = \chi.$$

$$S^{(S)} = S_E^{(S)} - \frac{\gamma M_P^2}{3} \int d^4x [\Delta(\Psi - \Phi)]^2$$

$$\frac{S_E^{(S)}}{M_P^2} = \frac{1}{2} \int d^4x a^2 \left[-6\Phi'^2 + 12\mathcal{H}\Psi\Phi' + 2\partial_i\Phi(2\partial^i\Psi + \partial^i\Phi) - 2(\mathcal{H}' + 2\mathcal{H}^2)\Psi^2 \right. \\ \left. + \frac{1}{M_P^2} (\chi'^2 - \partial_i\chi\partial^i\chi - a^2 V_{,\phi\phi}\chi^2 - 6\phi'\Phi'\chi - 2\phi'\chi'\Psi - 2a^2 V_{,\phi}\Psi\chi) \right]$$

For de Sitter background – **exact diagonalization**. Vacuum initial conditions on field perturbations at $\eta \rightarrow \infty$:

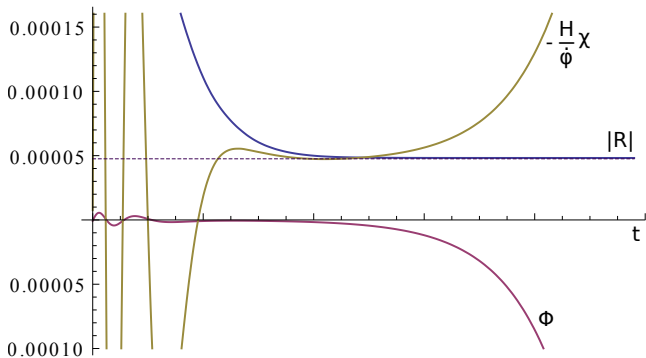
$$\Phi = \frac{H_0 q \eta}{2\pi \sqrt{6} M_P} e^{-iq\eta}, \quad \chi = \frac{H_0}{2\pi} q \eta e^{-iq\eta}$$

Superhorizon solutions

$$\Phi \sim a, \quad \chi \sim a, \quad \mathcal{R} = \Phi - \frac{H}{\dot{\phi}} \chi = -H_0 / (2\pi\sqrt{2\epsilon}M_p)$$

Superhorizon solutions

$$\Phi \sim a, \quad \chi \sim a, \quad \mathcal{R} = \Phi - \frac{H}{\dot{\phi}} \chi = -H_0 / (2\pi\sqrt{2\epsilon}M_p)$$



The comoving gauge

The Einstein equations may be written as one 4-th order equation on $W = \Psi - \Phi$. The four solutions in the **superhorizon limit** and for $\gamma H^2 \gg 1$ are:

The comoving gauge

The Einstein equations may be written as one 4-th order equation on $W = \Psi - \Phi$. The four solutions in the **superhorizon limit** and for $\gamma H^2 \gg 1$ are:

$$W_c,$$

$$\int a dt,$$

$$\int \frac{dt}{a},$$

$$\int a dt_1 \int^{t_1} \frac{H}{a^2} dt_2 \int^{t_2} \frac{a \dot{H}}{H^2} dt_3 \propto (t, t^2)$$

The comoving gauge

The Einstein equations may be written as one 4-th order equation on $W = \Psi - \Phi$. The four solutions in the **superhorizon limit** and for $\gamma H^2 \gg 1$ are:

$$W_c, \quad \int a dt, \quad \int \frac{dt}{a}, \quad \int a dt_1 \int^{t_1} \frac{H}{a^2} dt_2 \int^{t_2} \frac{a \dot{H}}{H^2} dt_3 \propto (t, t^2)$$

The Newtonian gauge breaks: $\Psi, \Phi \sim a \Rightarrow \Psi, \Phi > 1$.

The comoving gauge

The Einstein equations may be written as one 4-th order equation on $W = \Psi - \Phi$. The four solutions in the **superhorizon limit** and for $\gamma H^2 \gg 1$ are:

$$W_c, \quad \int a dt, \quad \int \frac{dt}{a}, \quad \int a dt_1 \int^{t_1} \frac{H}{a^2} dt_2 \int^{t_2} \frac{a \dot{H}}{H^2} dt_3 \propto (t, t^2)$$

The Newtonian gauge breaks: $\Psi, \Phi \sim a \Rightarrow \Psi, \Phi > 1$.

The comoving gauge:

$$g_{00} = a^2(1 + 2A), \quad g_{0i} = 2a^2 \partial_i B, \quad g_{ij} = a^2(\delta_{ij}(1 + 2\mathcal{R}) + \partial_i \partial_j E)$$

The comoving gauge

The Einstein equations may be written as one 4-th order equation on $W = \Psi - \Phi$. The four solutions in the **superhorizon limit** and for $\gamma H^2 \gg 1$ are:

$$W_c, \quad \int a dt, \quad \int \frac{dt}{a}, \quad \int a dt_1 \int^{t_1} \frac{H}{a^2} dt_2 \int^{t_2} \frac{a \dot{H}}{H^2} dt_3 \propto (t, t^2)$$

The Newtonian gauge breaks: $\Psi, \Phi \sim a \Rightarrow \Psi, \Phi > 1$.

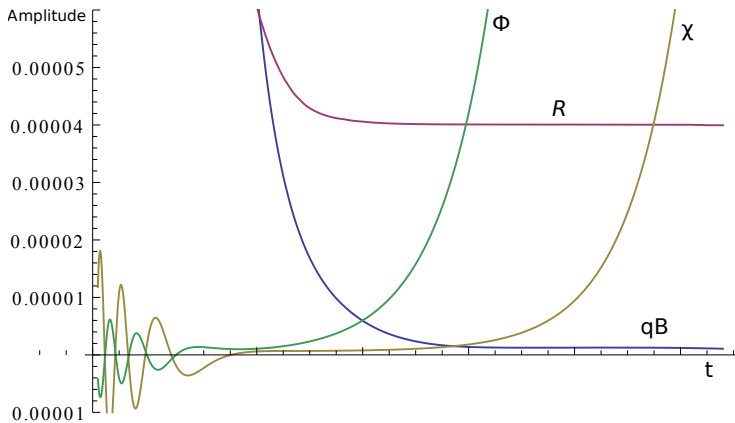
The comoving gauge:

$$g_{00} = a^2(1 + 2A), \quad g_{0i} = 2a^2 \partial_i B, \quad g_{ij} = a^2(\delta_{ij}(1 + 2\mathcal{R}) + \partial_i \partial_j E)$$

$$E = 0, \quad \mathcal{R} = -\frac{W_c}{2\epsilon} = -\frac{H_0}{2\pi\sqrt{2\epsilon}M_P},$$

$$B = \frac{\chi}{a\dot{\phi}} = \sqrt{\frac{\epsilon}{3q}} \mathcal{R}$$

Picture



Ghost becomes heavier than the Hubble parameter:

$$\gamma H^2 \ll 1$$

$$a \propto t^p, \quad p < 1$$

Ghost becomes heavier than the Hubble parameter: $\gamma H^2 \ll 1$

$$a \propto t^p, \quad p < 1$$

Superhorizon solutions for W :

$$W_c, \quad t^{-p-1}, \quad t^{p/2} e^{\pm it/\sqrt{\gamma}}$$

Ghost becomes heavier than the Hubble parameter: $\gamma H^2 \ll 1$

$$a \propto t^p, \quad p < 1$$

Superhorizon solutions for W :

$$W_c, \quad t^{-p-1}, \quad t^{p/2} e^{\pm it/\sqrt{\gamma}}$$

Sewing this solutions to previous limit use two important properties:

- Solutions $t^{p/2} e^{\pm it/\sqrt{\gamma}}$ do not impact on \mathcal{R}
- Einstein solutions W_c, t^{-p-1} do not impact on B

Ghost becomes heavier than the Hubble parameter:

$$\gamma H^2 \ll 1$$

$$a \propto t^p, \quad p < 1$$

Superhorizon solutions for W :

$$W_c, \quad t^{-p-1}, \quad t^{p/2} e^{\pm it/\sqrt{\gamma}}$$

Sewing this solutions to previous limit use two important properties:

- Solutions $t^{p/2} e^{\pm it/\sqrt{\gamma}}$ do not impact on \mathcal{R}
- Einstein solutions W_c, t^{-p-1} do not impact on B

Result:

$$\mathcal{R} = const, \quad B = \sqrt{\frac{\epsilon}{3}} \frac{\mathcal{R}}{q} \left(\frac{t}{\sqrt{\gamma}} \right)^{-p/2} e^{\pm it/\sqrt{\gamma}}$$

Crossing back the horizon

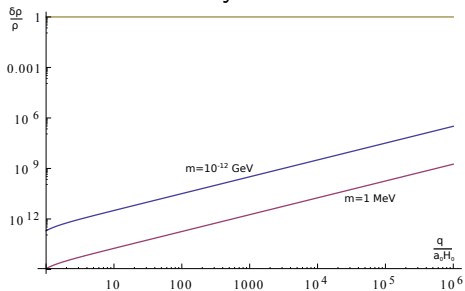
- The ghost solutions $W \sim t^{p/2} e^{\pm it/\sqrt{\gamma}}$ remains unchanged after horizon crossing.

Crossing back the horizon

- The ghost solutions $W \sim t^{p/2} e^{\pm it/\sqrt{\gamma}}$ remains unchanged after horizon crossing.
- Einstein modes behaves as usual.

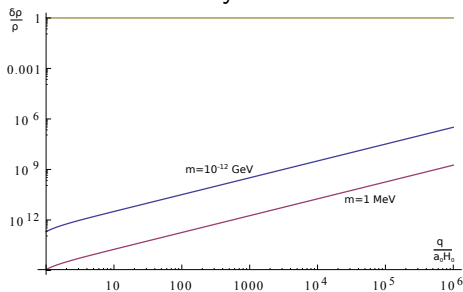
Crossing back the horizon

- The ghost solutions $W \sim t^{p/2} e^{\pm it/\sqrt{\gamma}}$ remains unchanged after horizon crossing.
- Einstein modes behaves as usual.
- Ghost impact to the matter density contrast:



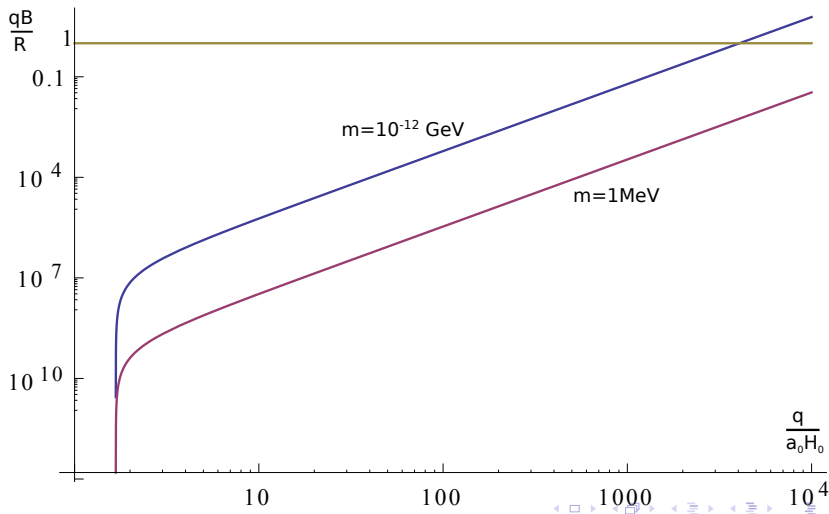
Crossing back the horizon

- The ghost solutions $W \sim t^{p/2} e^{\pm it/\sqrt{\gamma}}$ remains unchanged after horizon crossing.
- Einstein modes behaves as usual.
- Ghost impact to the matter density contrast:

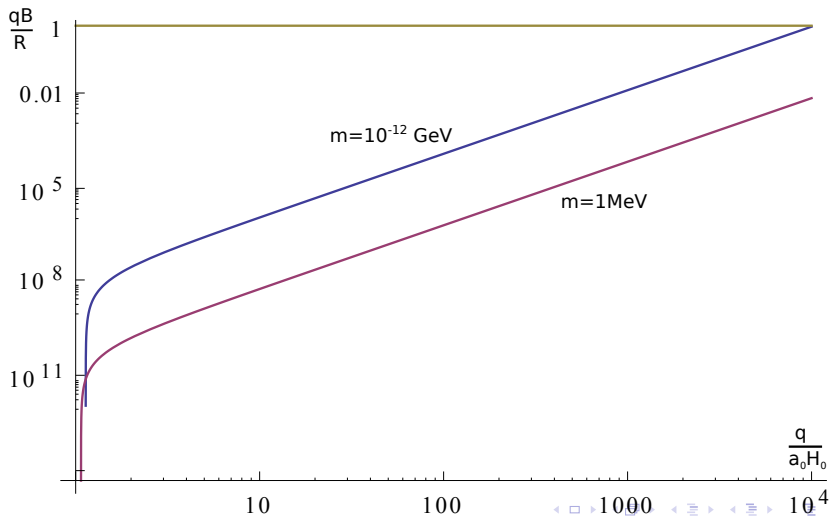


- The relative ghost impact to CMB is characterized by $g_{0i} = n_i q B / \mathcal{R}$ in comoving gauge.

$n_{ig_{oi}}$ spectrum at recombination



n_{igoi} spectrum now



Conclusions

- It looks possible to use the renormalizable quadratic gravity for constructing inflationary models with controlled quantum corrections

Conclusions

- It looks possible to use the renormalizable quadratic gravity for constructing inflationary models with controlled quantum corrections
- The light ghost limit correspond to very weak coupling to matter:

Conclusions

- It looks possible to use the renormalizable quadratic gravity for constructing inflationary models with controlled quantum corrections
- The light ghost limit correspond to very weak coupling to matter:
 - Negligible production in a hot plasma
 - Lifetime is much larger than the Universe lifetime

Conclusions

- It looks possible to use the renormalizable quadratic gravity for constructing inflationary models with controlled quantum corrections
- The light ghost limit correspond to very weak coupling to matter:
 - Negligible production in a hot plasma
 - Lifetime is much larger than the Universe lifetime

$$10^{-3} \text{ eV} \lesssim m \lesssim 10 \text{ MeV}$$

Conclusions

- It looks possible to use the renormalizable quadratic gravity for constructing inflationary models with controlled quantum corrections
- The light ghost limit correspond to very weak coupling to matter:
 - Negligible production in a hot plasma
 - Lifetime is much larger than the Universe lifetime

$$\downarrow$$
$$10^{-3} \text{ eV} \lesssim m \lesssim 10 \text{ MeV}$$

- Are very small negative probabilities dangerous?

Conclusions

- It looks possible to use the renormalizable quadratic gravity for constructing inflationary models with controlled quantum corrections
- The light ghost limit correspond to very weak coupling to matter:
 - Negligible production in a hot plasma
 - Lifetime is much larger than the Universe lifetime

$$\downarrow$$
$$10^{-3} \text{ eV} \lesssim m \lesssim 10 \text{ MeV}$$

- **Are very small negative probabilities dangerous?**
- The cosmological ghost scalar perturbations do not affect the Planck data.

Thanks for your attention!